ROTATIONAL TRANSITIONS IN THE INTERACTING ν_2 , ν_3 , ν_4 AND ν_6 BANDS OF FORMALDEHYDE IN H₂¹²C¹⁶O THE MILLIMETER RANGE FOR ASTROPHYSICAL USE

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This work, besides its fundamental interest, is motivated by the astrophysical importance of formaldehyde. For example formaldehyde was detected by millimeter techniques in Orion-KL and in several low-mass protostars ^{*a*}. However, no line parameters are presently available in the spectroscopic databases for the rotational transitions within the 2¹, 3¹, 4¹ and 6¹ first excited vibrational states of formaldehyde. The goal of this study is to generate a list of line parameters for these "hot" transitions in order to help such - may be - future identifications in astrophysical spectra. For this reason, submillimeter spectra were recorded at Lille and at Köln in the 150-650 and 850-900 GHz spectral ranges, respectively. These sub millimeter data were combined in a least squares fit calculation with the infrared experimental data available in the literature for the ν_3 , ν_4 and ν_6 bands ^{*b*}, and for the ν_2 band ^{*c*}. The Hamiltonian model accounts for the various Coriolis-type resonances which perturb the energy levels of the 3¹, 4¹ and 6¹ vibrational states. In addition a weaker and somehow unexpected anharmonic resonance coupling the 2¹ and 3¹ energy levels was accounted for. Using this theoretical model, it proved possible to reproduce satisfactorily the experimental data and to generate a list of line positions and intensities for the $\nu_2 \leftrightarrow \nu_2$, $\nu_3 \leftrightarrow \nu_3$, $\nu_4 \leftrightarrow \nu_4$, and $\nu_6 \leftrightarrow \nu_6$ rotational transitions.

^aMangum-JG; Wootten-A; Barsony-M, Astrophys. J. 526 845-53 (1999)

^bD. C. Reuter, H. Takeo, S. Nadler, S.J.Daunt, and J. W. C. Johns, J. Chem. Phys. 91, 646 (1989)

^cF. Kwabia Tchana, A. Perrin and N. Lacome, J. Mol. Spectrosc. 245, 141-144, (2007)